Initial Results from the First Mass Loss Survey of Gas Giants Orbiting F Stars

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Close-in Planets Can Undergo Extreme Photoevaporation

XUV radiation



Hydrodynamic Outflow



Image Credit: NASA/CXC/M.Weiss



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An escaping atmos. should increase transit depth

Brightness

Broadband planet size



An escaping atmos. should increase transit depth

ghtness Briç

Broadband planet size

Helium 1083 nm planet size





An escaping atmos. should increase transit depth

ghtness Briç

Broadband planet size

Excess Absorption

Helium 1083 nm planet size





Metastable Helium outflow detections are widely successful but preferential to cooler stars



Confirmed Exoplanets Helium Outflow Detections M stars K stars G stars F stars

Allart et al. (2018, 2019); Mansfield et al. (2018); Nortmann et al. (2018); Oklopčić & Hirata (2018); Salz et al. (2018); Spake et al. (2018, 2021); Alonso-Floriano et al. (2019); Guilluy et al. (2020); Kirk et al. (2020, 2022); Palle et al. (2020); Ninan et al. (2020); Vissapragada et al. (2020, 2022); Paragas et al. (2021); Czesla et al. (2022); Zhang et al. (2022a,b); Gully-Santiago et al. (2023); Orell-Miquel et al. (2023); Pérez González et al. (2023); Zhang et al. (2023); Saidel et al. (2025)



First Two F star Detections Reveal Fastest Outflows







First Two F star Detections Reveal Fastest Outflows









Three potential drivers for the large outflows



High stellar rotational velocities

e.g. Johnstone et al. (2021)



Three potential drivers for the large outflows

Roche Lobe

High stellar rotational velocities



Planet fills large fraction of Roche lobe

Zhang et al. (2023); Gully-Santiago et al. (2023)



Three potential drivers for the large outflows

Roche Lobe

High stellar rotational velocities

Planet fills large fraction of Roche lobe



Additional Energy Sources (Balmer Escape?)*

*and/or different outflow geometries

García Muñoz & Schneider (2019)





















Survey F star worlds with 1083 nm Helium Filter on Palomar



Image Credit: Caltech/Palomar

Vissapragada et al. (2020)





SNR = 0.5SNR = 2.0HAT-P-32 b SNR = 3.5

Predicted based on 1D **EUV-driven** outflow models

50 30 60 40 vsin i [km/s]

Vissapragada et al. (2020); Dos Santos et al. (2022)





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SNR = 0.5SNR = 2.0HAT-P-32 b SNR = 3.5

30 50 60 40 vsin i [km/s]





vsin i [km/s]





vsin i [km/s]



Three tentative detections of atmospheric escape!



Saidel et al. in prep.

WASP-180 A b (3.7σ) 1.02 1.01 .0 Relative Flux 0.99 0.98 0.97 2024/07/28 UT 0.96 2024/12/06 UT 0.01 Residual 0.00 -0.010.05 0.00 -0.10 -0.050.00 0.05 Time from Eclipse Center [d] Time from Eclipse Center [d]

Helium 1083nm planet size

Broadband planet size









No clear solitary driver of fast outflows



	•	SNR = 0.5
b		SNR = 2.0
AT-P-32 b		SNR = 3.5

WASP-93 b

60

WASP-180 A b

30 50 40 vsin i [km/s]



What's Next?

Finish survey







3D Models

Stellar XUV spectra



Image Credit: James Vaughan



