



WFIRST Imaging and
Ground-based Spectroscopy
*A Powerful Combination for
Near-field Cosmology*

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Science in our Own Backyard with WFIRST, Caltech, Pasadena, CA

Next Generation Virgo Cluster Survey

— *CFHT/MegaCam imaging (104 deg²)*
— *Keck/DEIMOS spectroscopy*

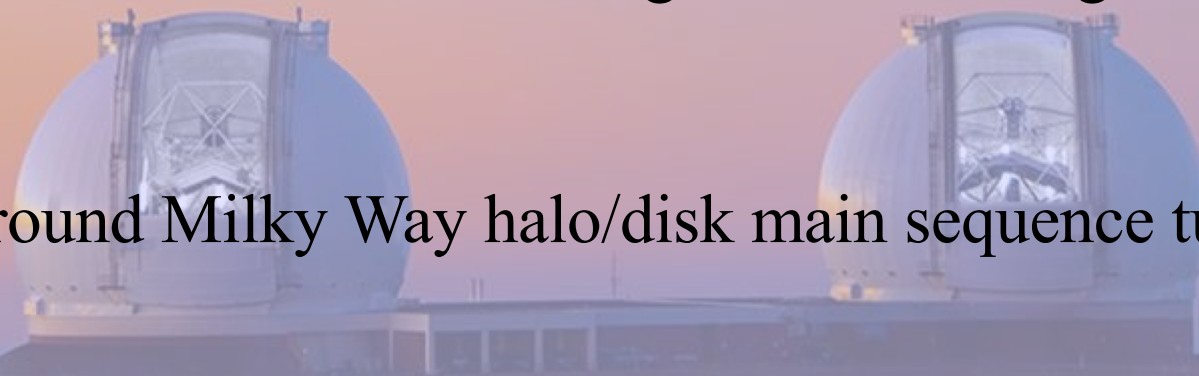
Collaborators:

Pat Côté, Yuting Feng, Laura Ferrarese, Biao Li, Sungsoon Lim, Chris Mihos,
Alessia Longobardi, Eric Peng, Elisa Toloba, Yunhao Zhang

*Stephanie Chen, Jason Chu, Rachel Guo, Teddy Liang, Tiffany Louie, Adhara
Martellini, Celia Ramirez, Arushi Sahai, Andrew Shao, Lea Sparkman,
Kovid Tallam, Jin Tuan*

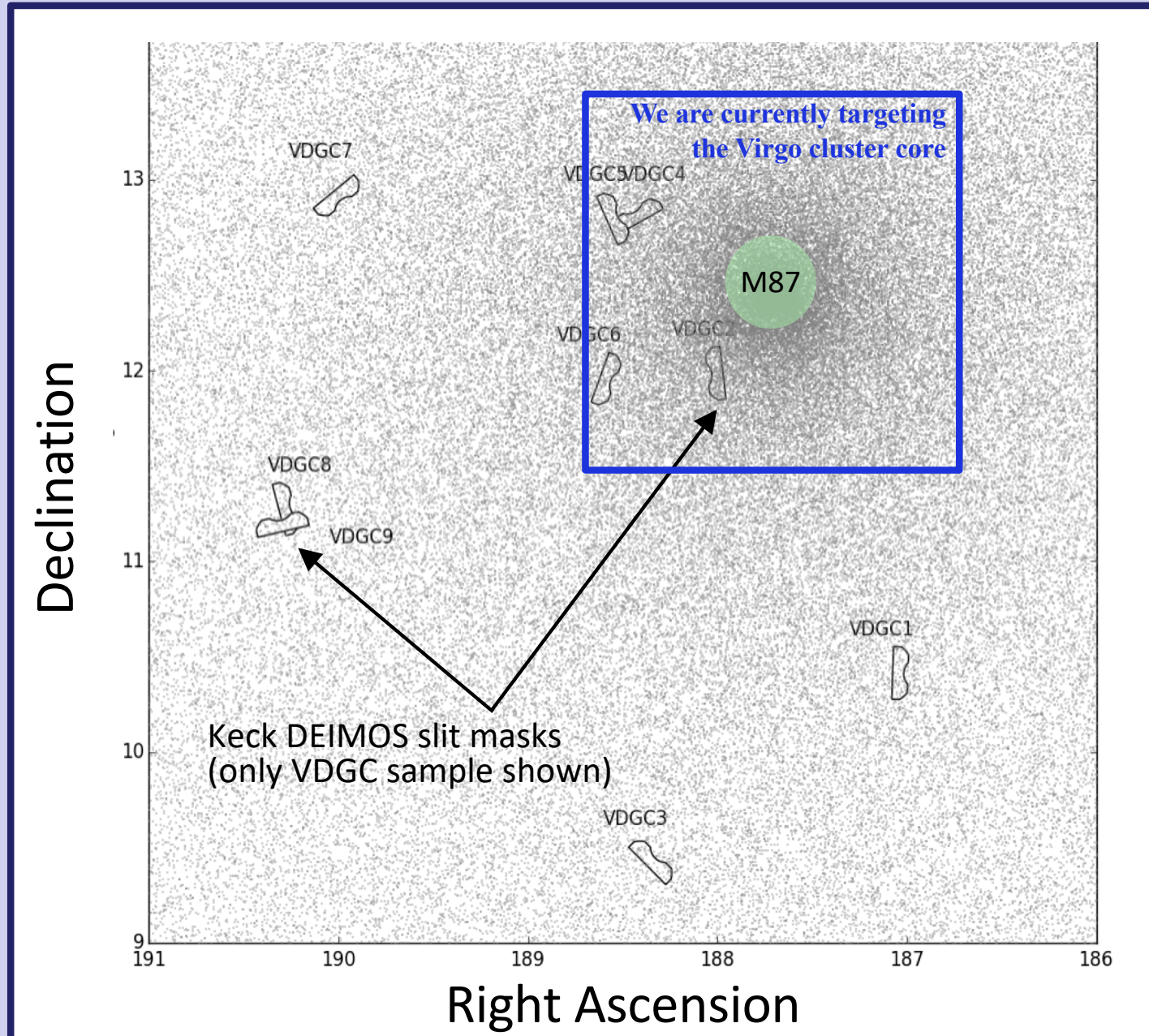
NGVS Keck/DEIMOS Spectroscopy

- ❖ Keck/DEIMOS multislit spectroscopy: 600 lines/mm grating, medium resolution ($R \sim 2000$), 4800–9500Å wavelength coverage
- ❖ Targets
 - Candidate GC satellites of Virgo dwarf elliptical galaxies (dEs)
 - Candidate GC satellites of Virgo ultra diffuse galaxies (UDGs)
- ❖ Fillers
 - Foreground Milky Way halo/disk main sequence turnoff stars

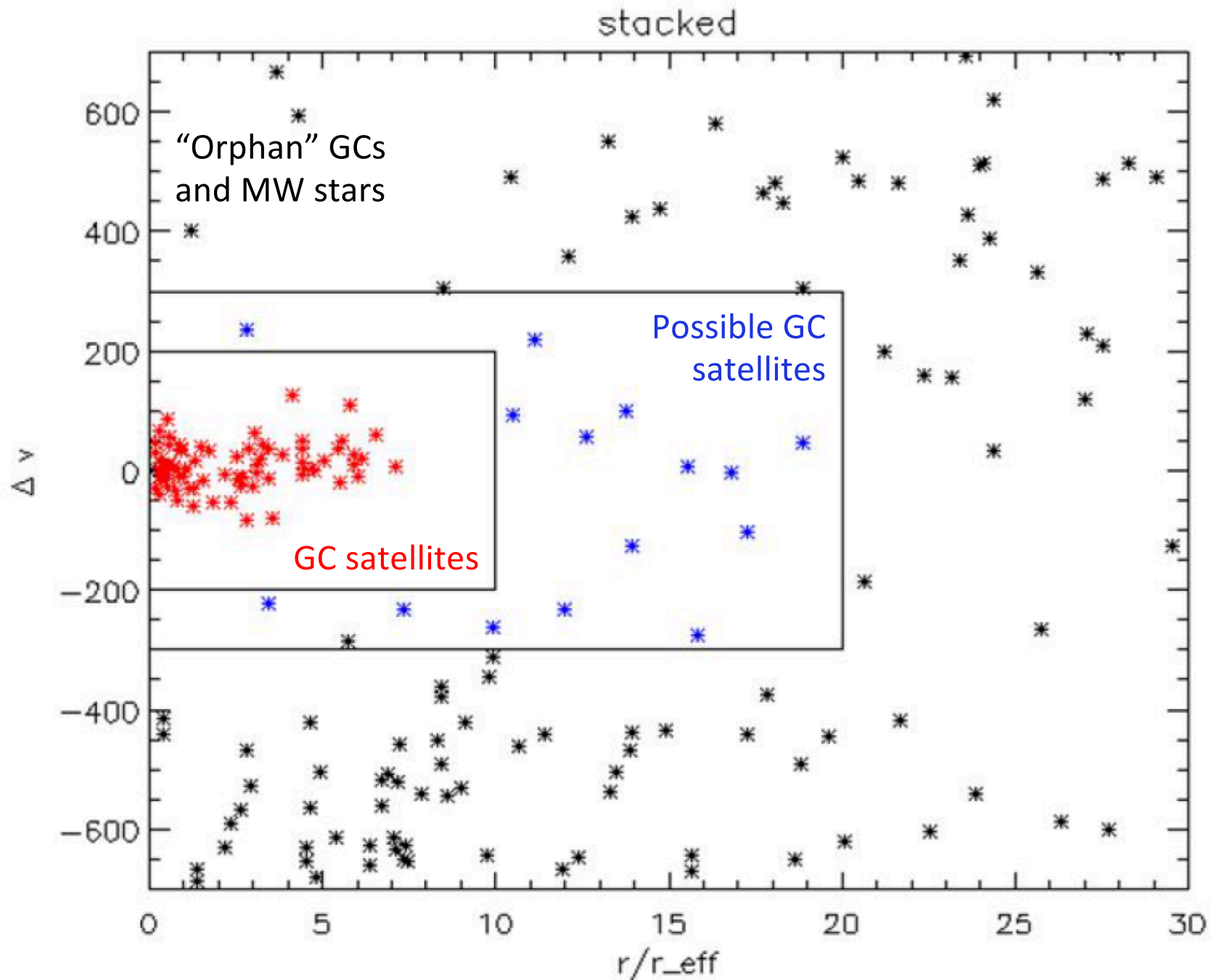


Keck/DEIMOS Spectroscopic Sample

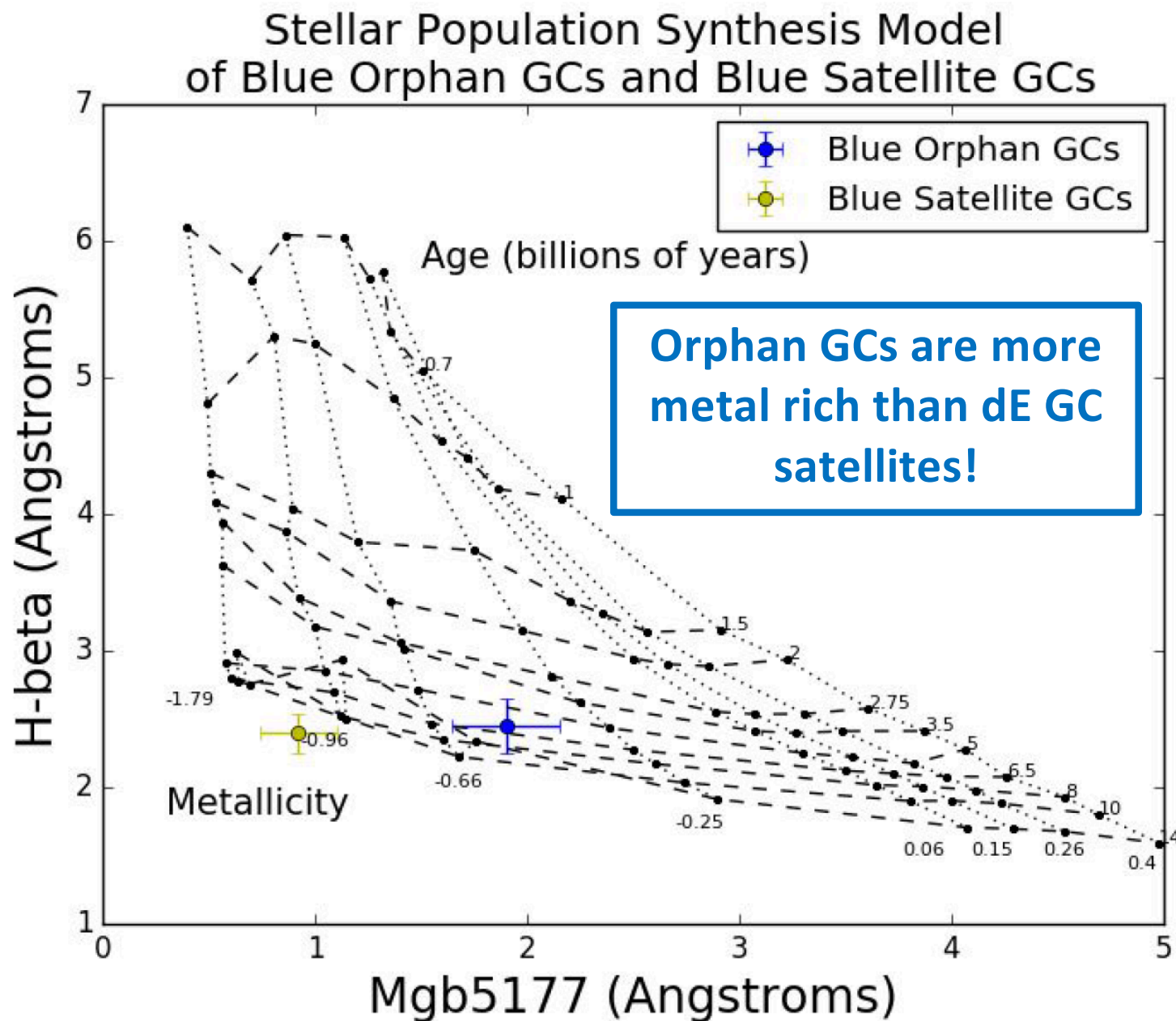
Keck/DEIMOS spectroscopy of GC satellites of Virgo cluster dEs



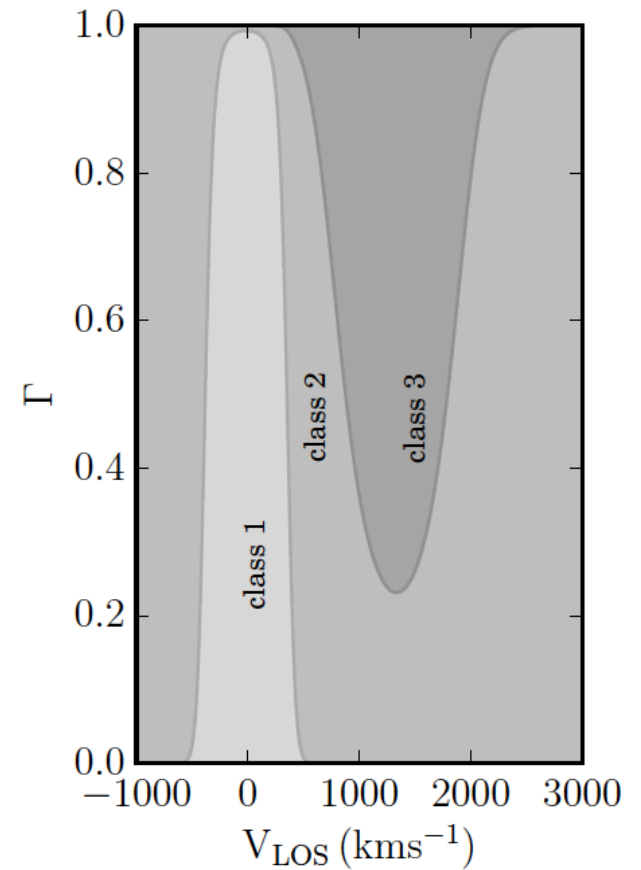
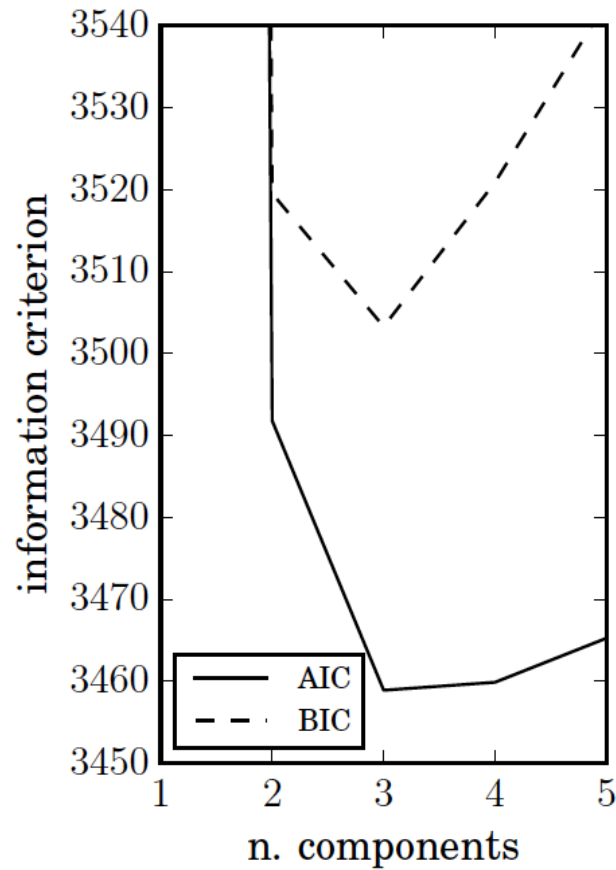
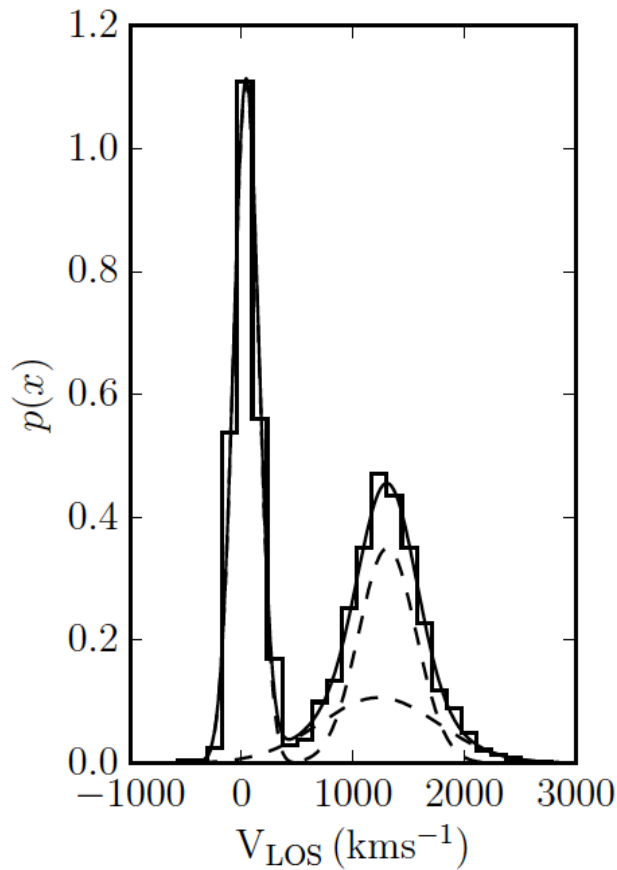
Definition of GC satellites versus “orphan” GCs



Comparing ages and metallicities of “orphan” GCs and dE GC satellites



Two flavors of “orphan” GCs: M87 halo GCs and intracluster GCs

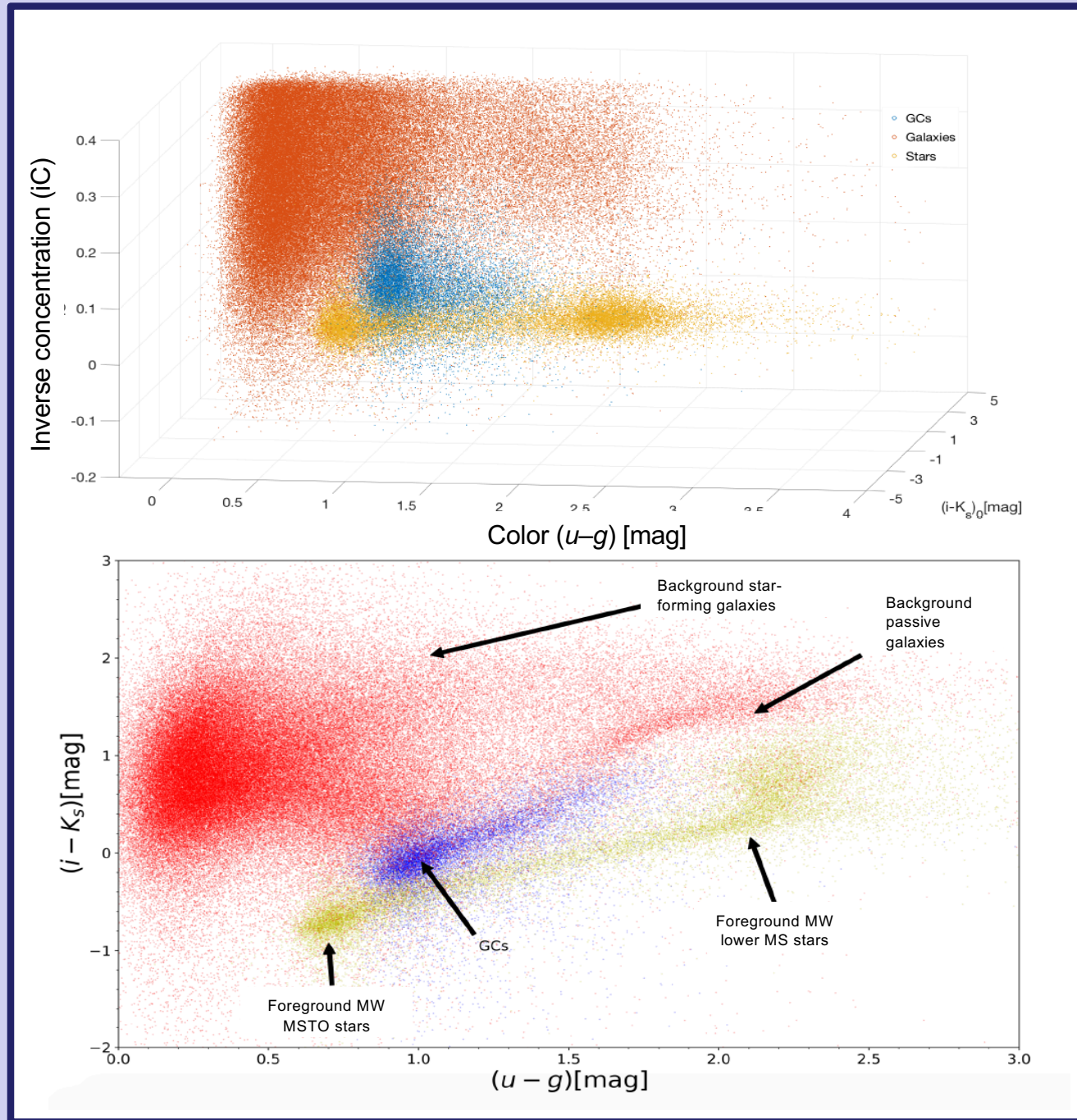


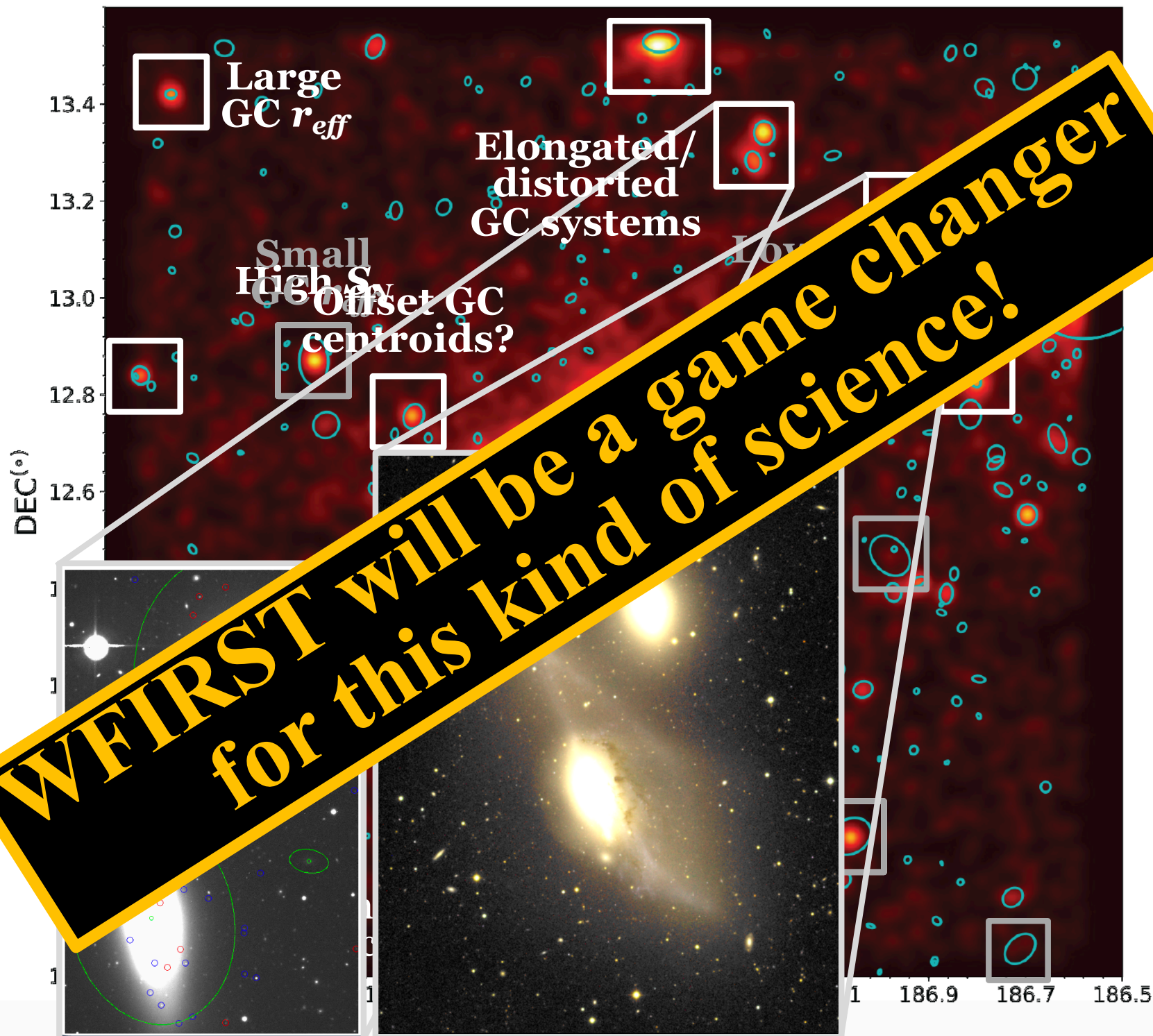
- M87 halo GCs: $\langle v \rangle = 1300 \text{ km/s}$; $\sigma_v \sim 400 \text{ km/s}$
- Intracluster GCs: $\langle v \rangle = 1100 \text{ km/s}$; $\sigma_v \sim 700 \text{ km/s}$

Virgo Core Photometric GC Sample

- Photometric sample of GCs selected using Eric Peng's "extreme deconvolution" (XD) method based on broadband colors ($ugizK_s$) and image morphology
- This has resulted in the most complete map to date of GCs in Virgo cluster core ($r < 200$ kpc)

GC selection based on “extreme deconvolution”





WFIRST will be a game changer for this kind of science!

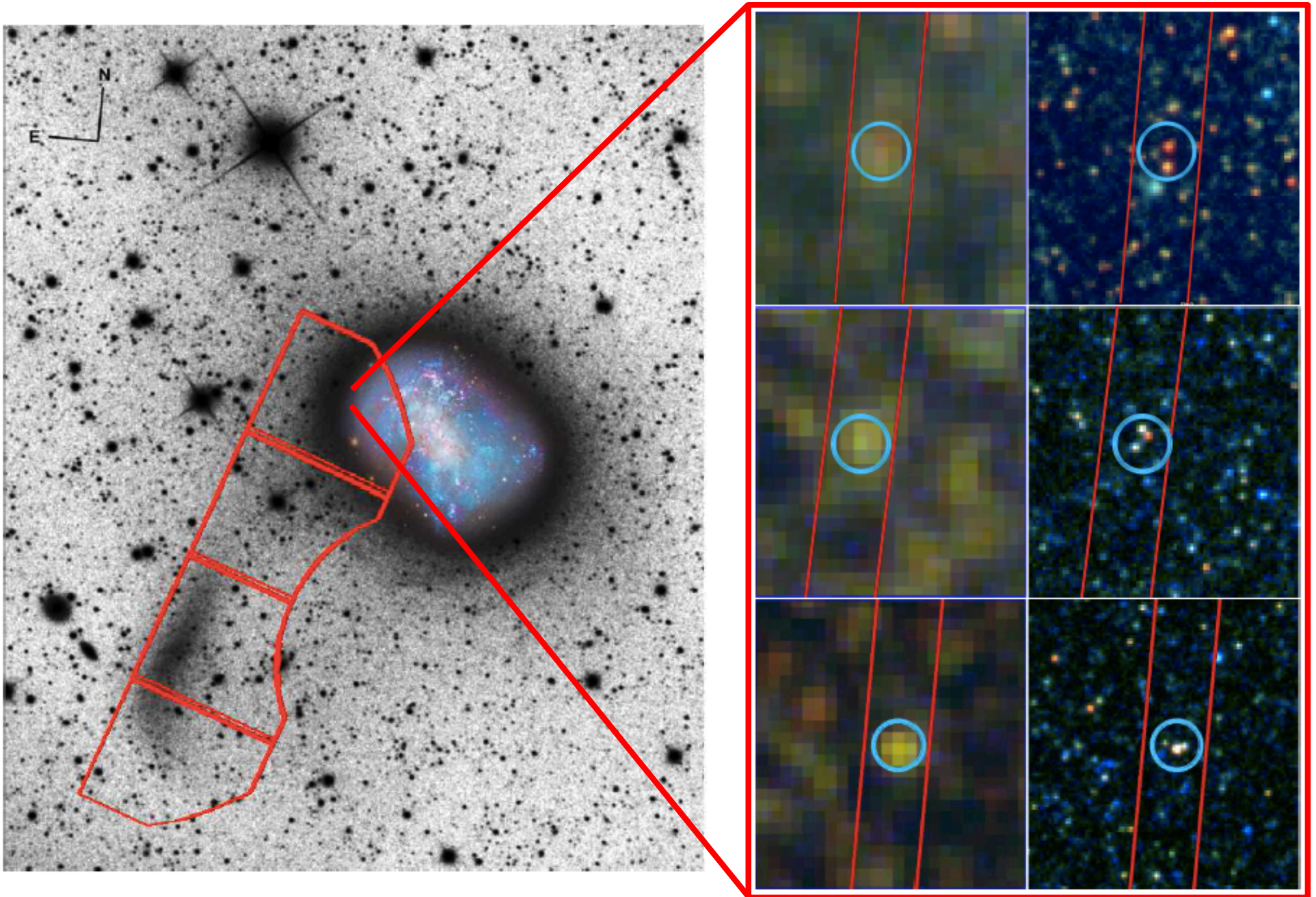
Panoramic Imaging Survey of Centaurus and Sculptor (PISCeS)

- *Magellan/MegaCam imaging*
- *Keck/DEIMOS spectroscopy*

Collaborators:

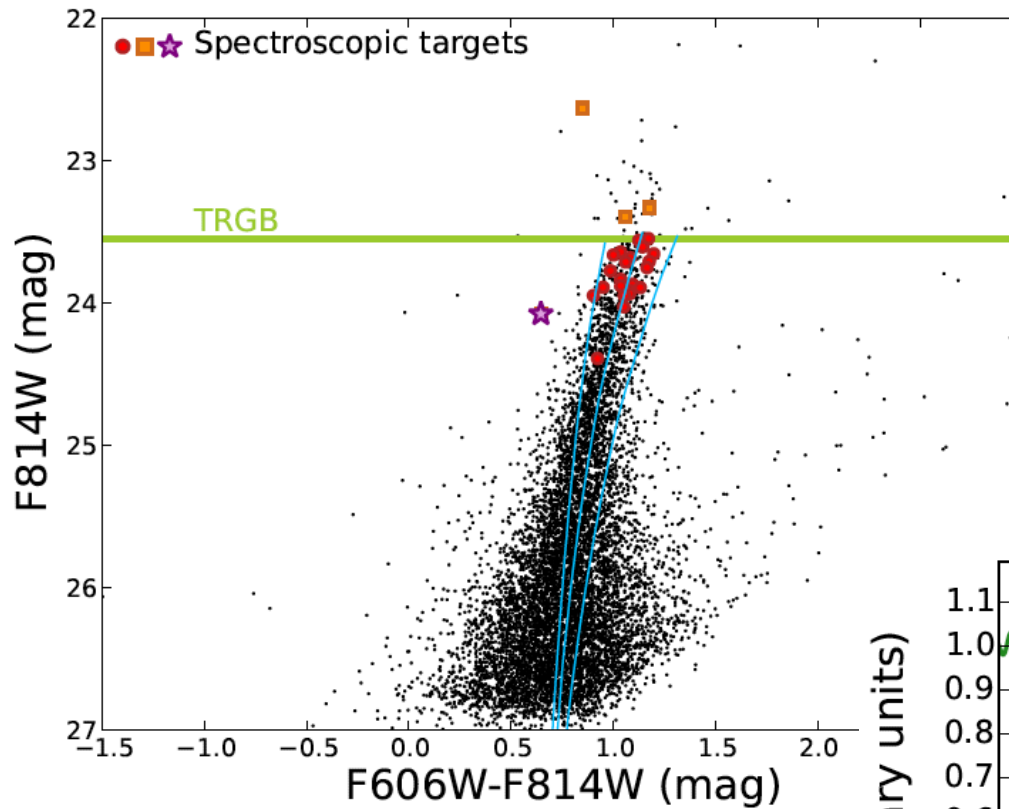
Denija Crnojević, Dave Sand, Josh Simon, Elisa Toloba

Coadded spectroscopy of positive SBFs in a LSB stellar stream



Toloba, PG, et al. (2016a)

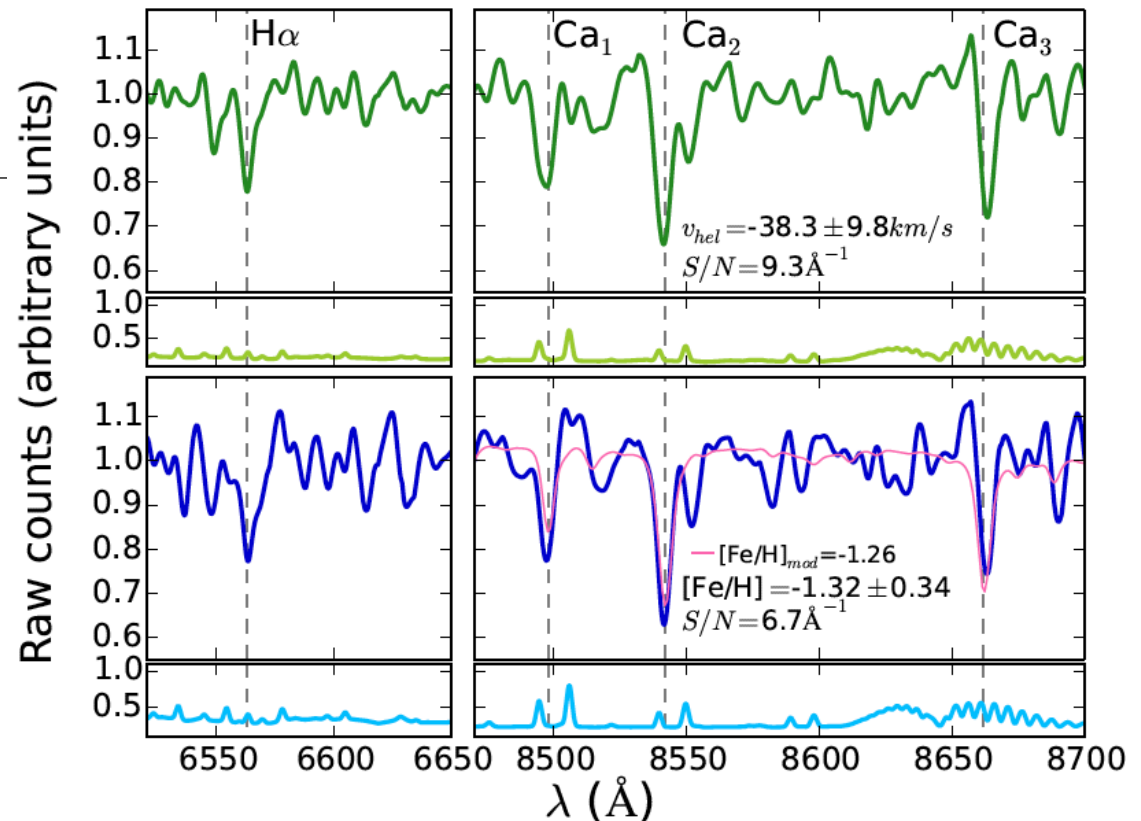
Coadded spectroscopy of stars & blends in an M81 dwarf satellite



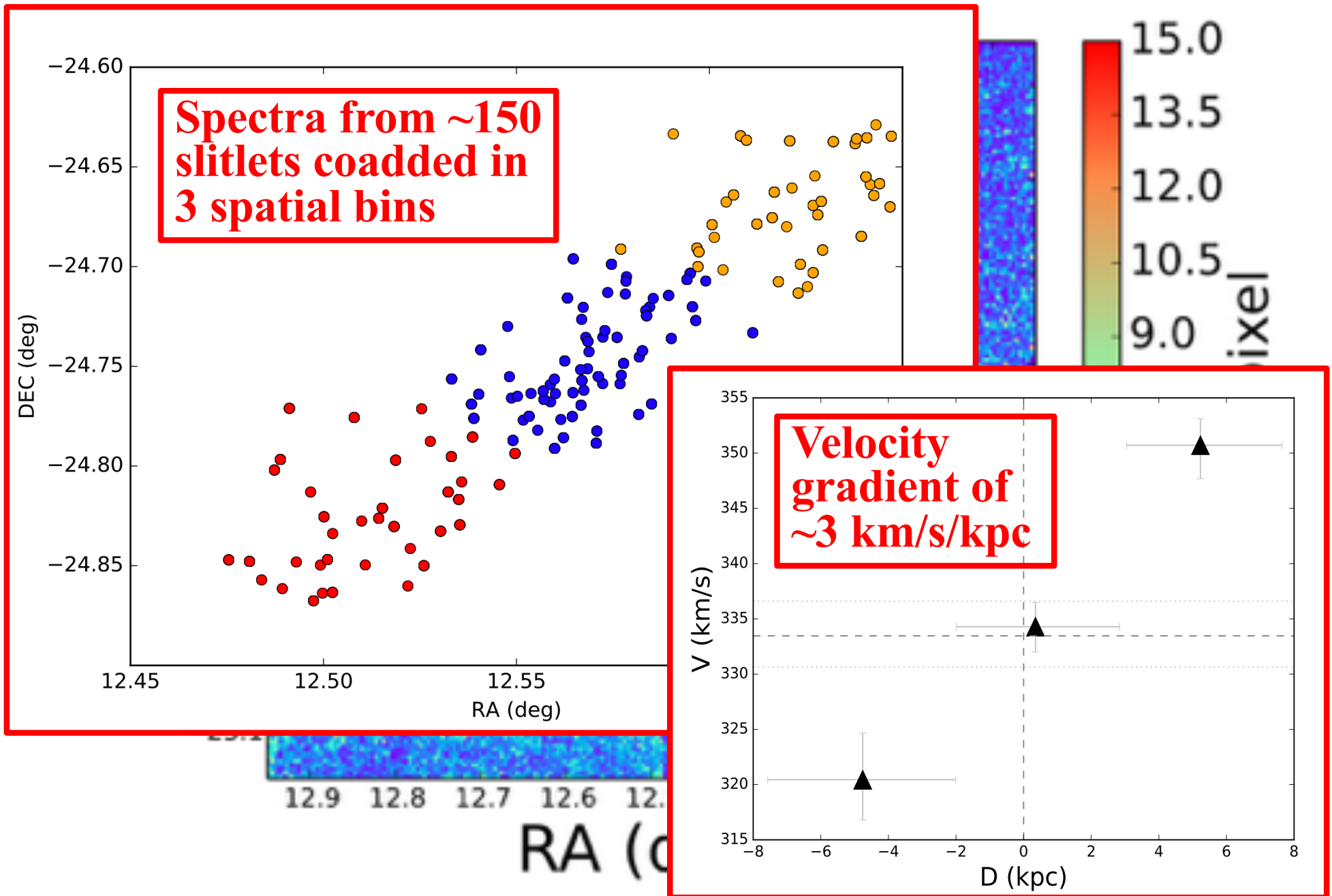
CMD of individual stars and blends in M81 dwarf satellite

Toloba et al (2016c)

Obvious H-alpha and Ca triplet absorption in the coadded spectrum



PISCeS survey: Tidal stream in NGC 253's halo (Sculptor group galaxy) with Keck/DEIMOS mask overlay



PISCeS survey: Tidal streams in the halo of Centaurus A



Elisa Toloba



Denija Crnojević

Sand+14, Crnojevic+14,
Crnojevic+16, Toloba+16

Accretion history and mass of the Milky Way

HALO7D survey

Looking at (and through) the remote Milky Way stellar halo



Emily Cunningham



Alis Deason

— Anderson, Cunningham, Deason, PG, Kirby, Rockosi, Sohn, van der Marel, Wetzel
— Barro, Conroy, Faber, Guo, Koo, Yesuf

“7D” mapping of the Milky Way halo: Accretion history and mass estimate

- ❖ Proper motions from multi-epoch imaging with HST, Gaia, and, in the future, WFIRST and JWST
- ❖ Keck/DEIMOS spectroscopy (8 to 32 hours per mask!) and ultimately spectroscopy with ELTs to measure radial velocities and especially chemical abundances of faint MSTO stars

HALO7D Collaboration

HSTPROMO: The HST Proper Motion Collaboration

Alis Deason, Emily Cunningham, Connie Rockosi, PG (UCSC), Evan Kirby (Caltech)
Roeland van der Marel, Jay Anderson, Tony Sohn (STScI)

Summary: Local Volume Science with WFIRST and Ground-based Spectroscopy

1. *Globular clusters as probes of the dark matter halo, star formation history, and chemical abundance of ultra-diffuse galaxies and dwarf ellipticals in the dense Virgo cluster environment*
2. *Coadded SBF spectroscopy of partially resolved stellar populations as a probe of the dark matter halos of galaxies out to distances of ~10 Mpc*
3. *6D phase space measurements of remote Milky Way halo stars and 3D kinematics of M31 and M33 resolved stellar populations*